

Forward modeling of coronal mass ejection flux ropes in the inner heliosphere with 3DCORE

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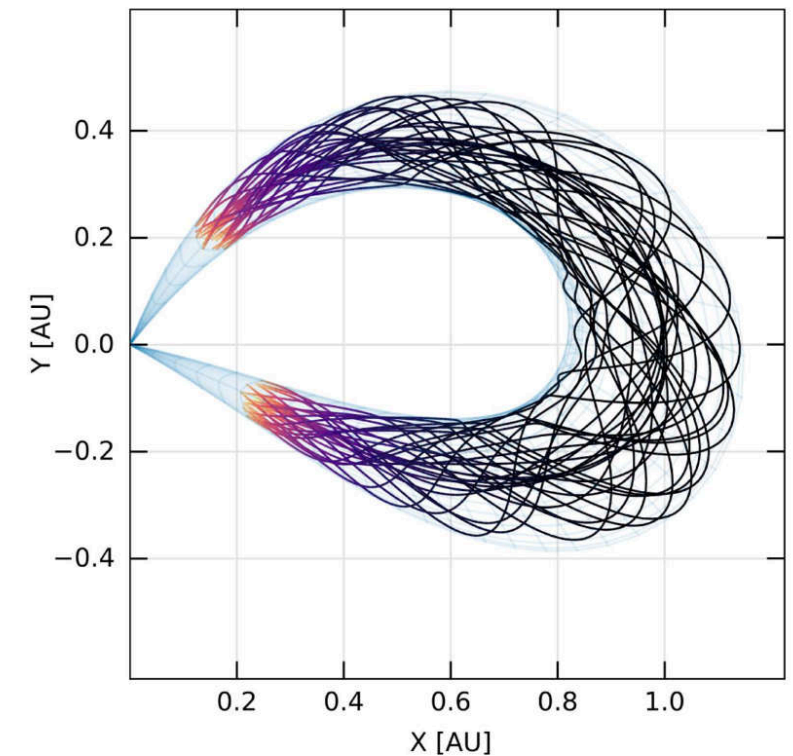
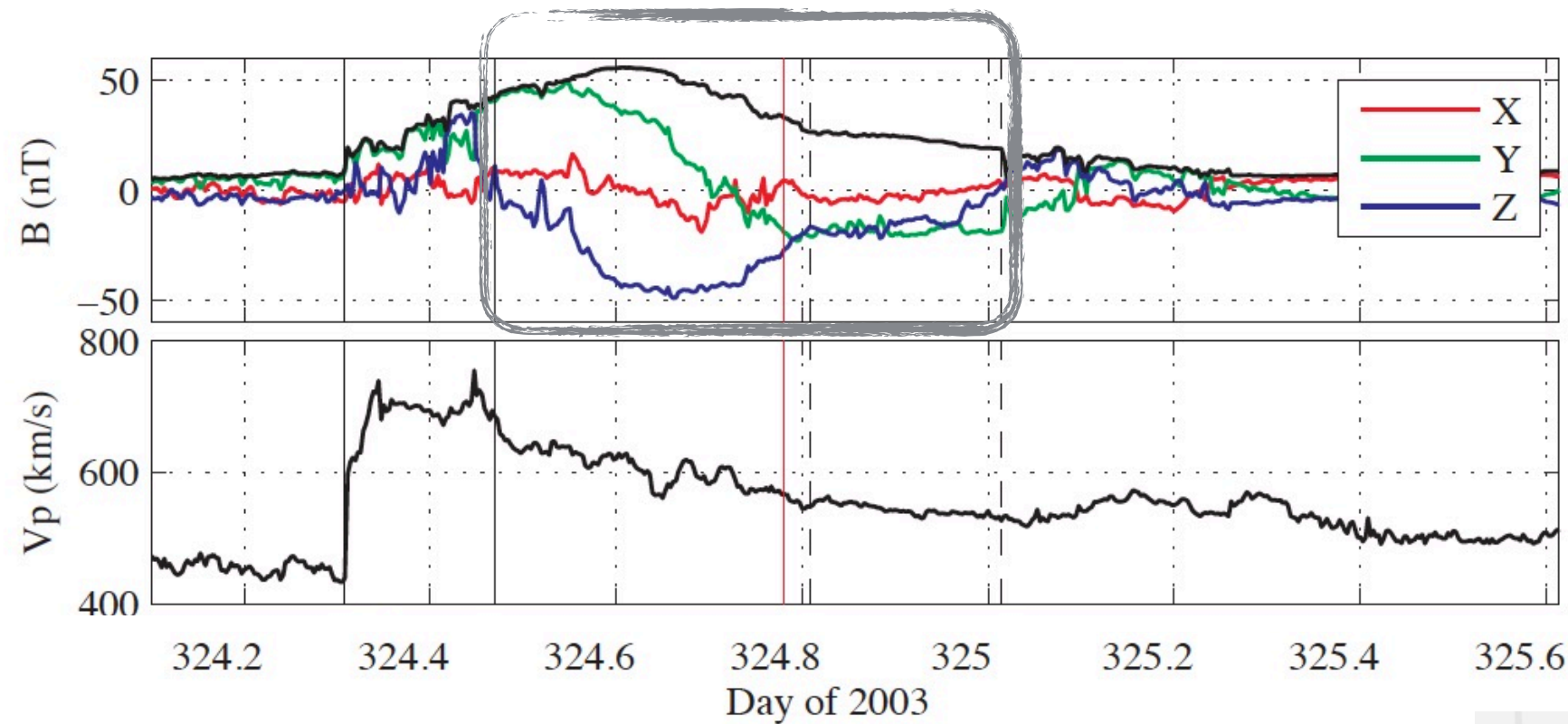
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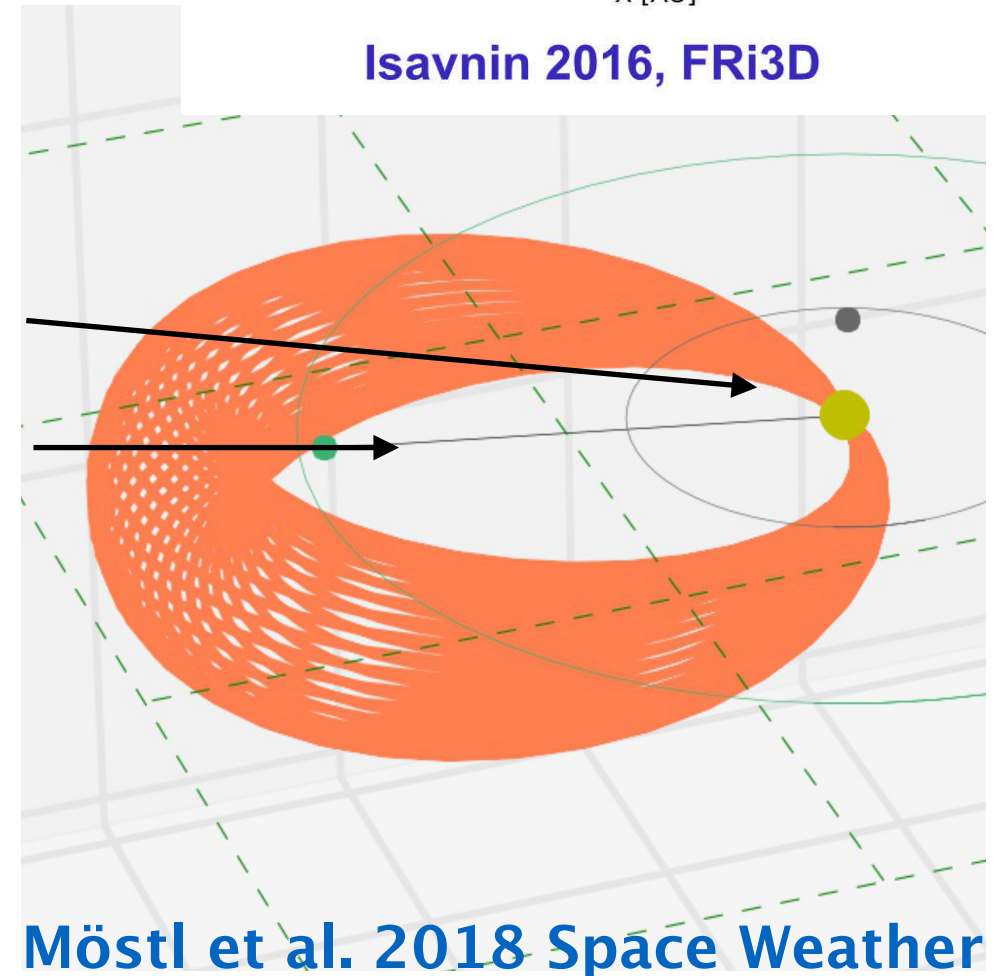
Bz problem



Semi empirical CME flux rope models:

- able to make magnetic field predictions
- 3D trajectory effects
- very quick computation times
- start with outer coronagraph field of view

Isavnin 2016, FRi3D

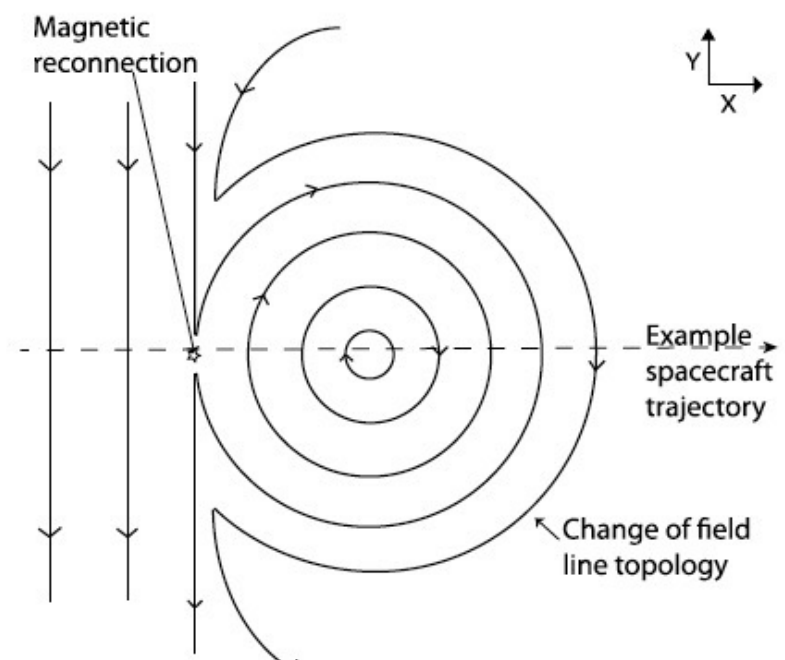
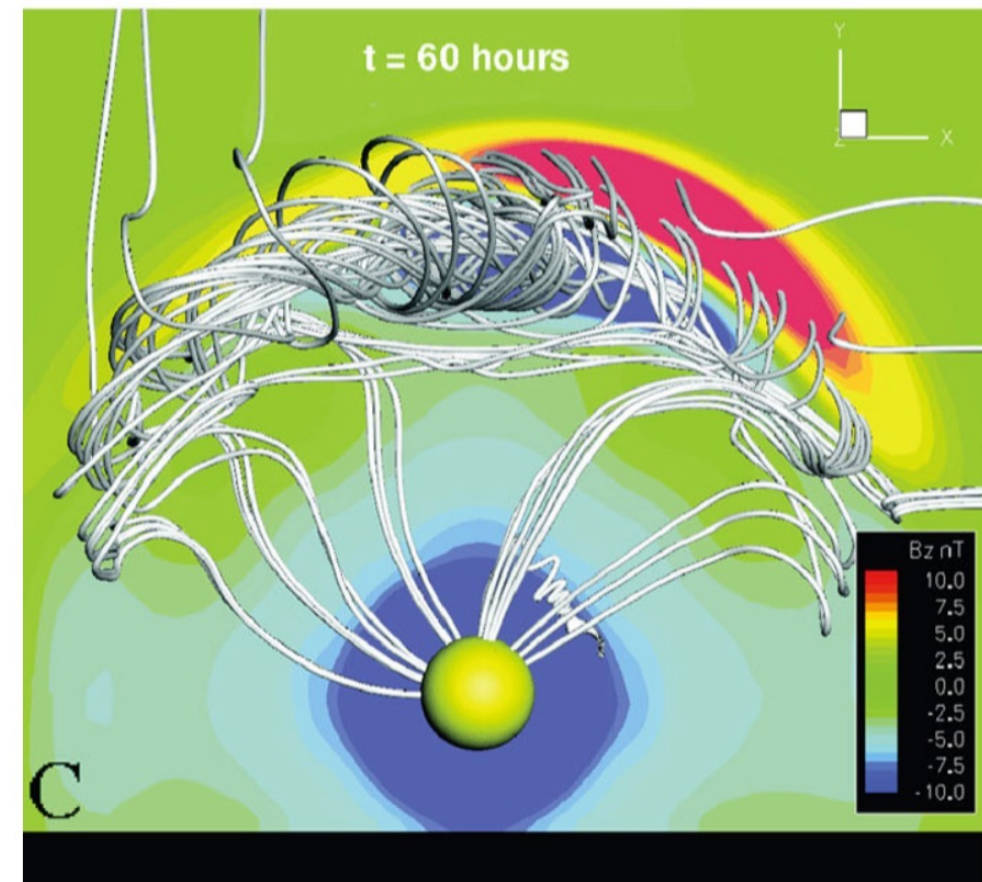


ICME physics

- Included:
- Deflection, channeling in the corona (IP unclear)
- Rotation in the corona (IP unclear)
- Expansion, decline of magnetic field
- Solar wind drag
- Pre-conditioning of solar wind

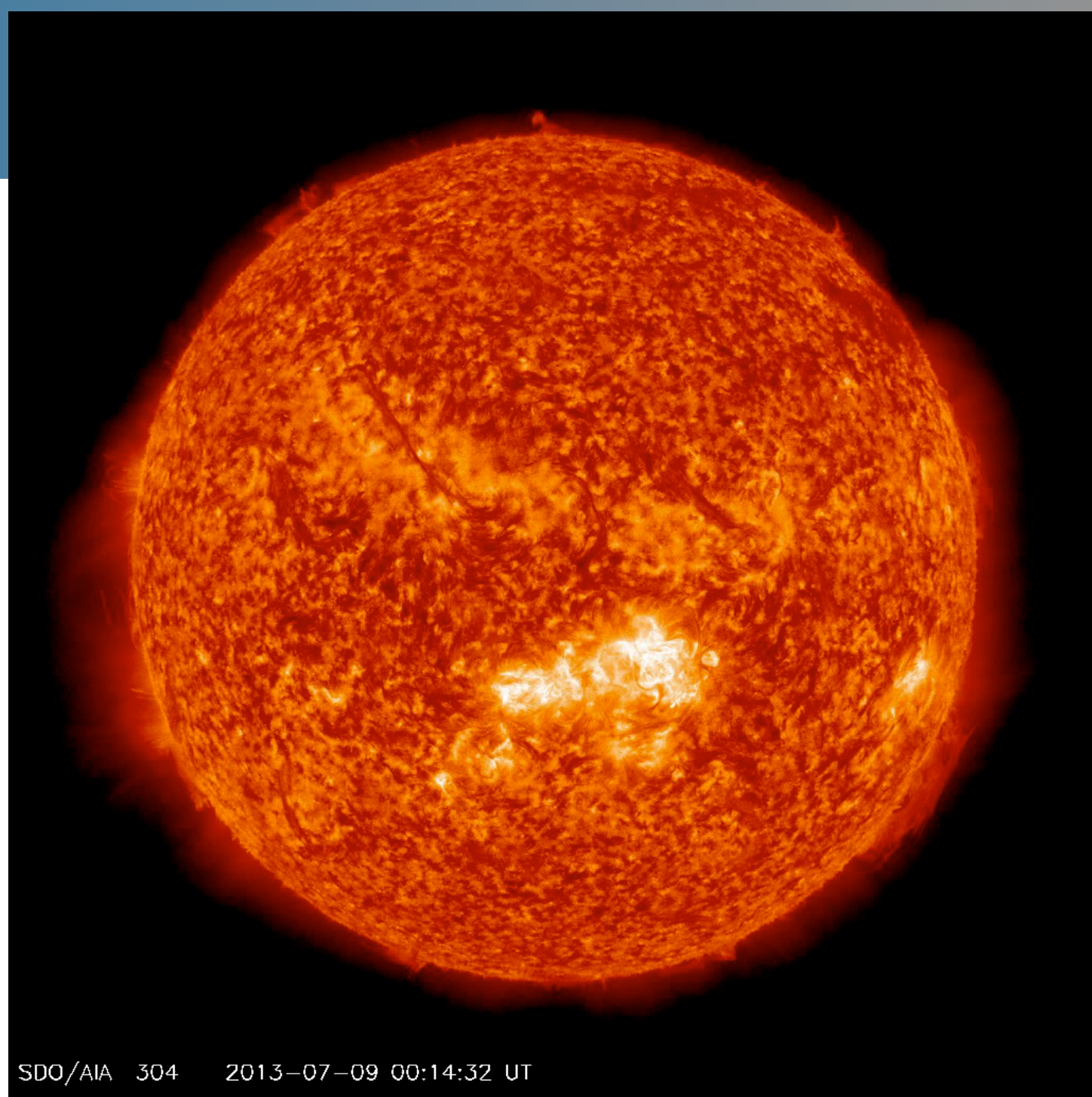
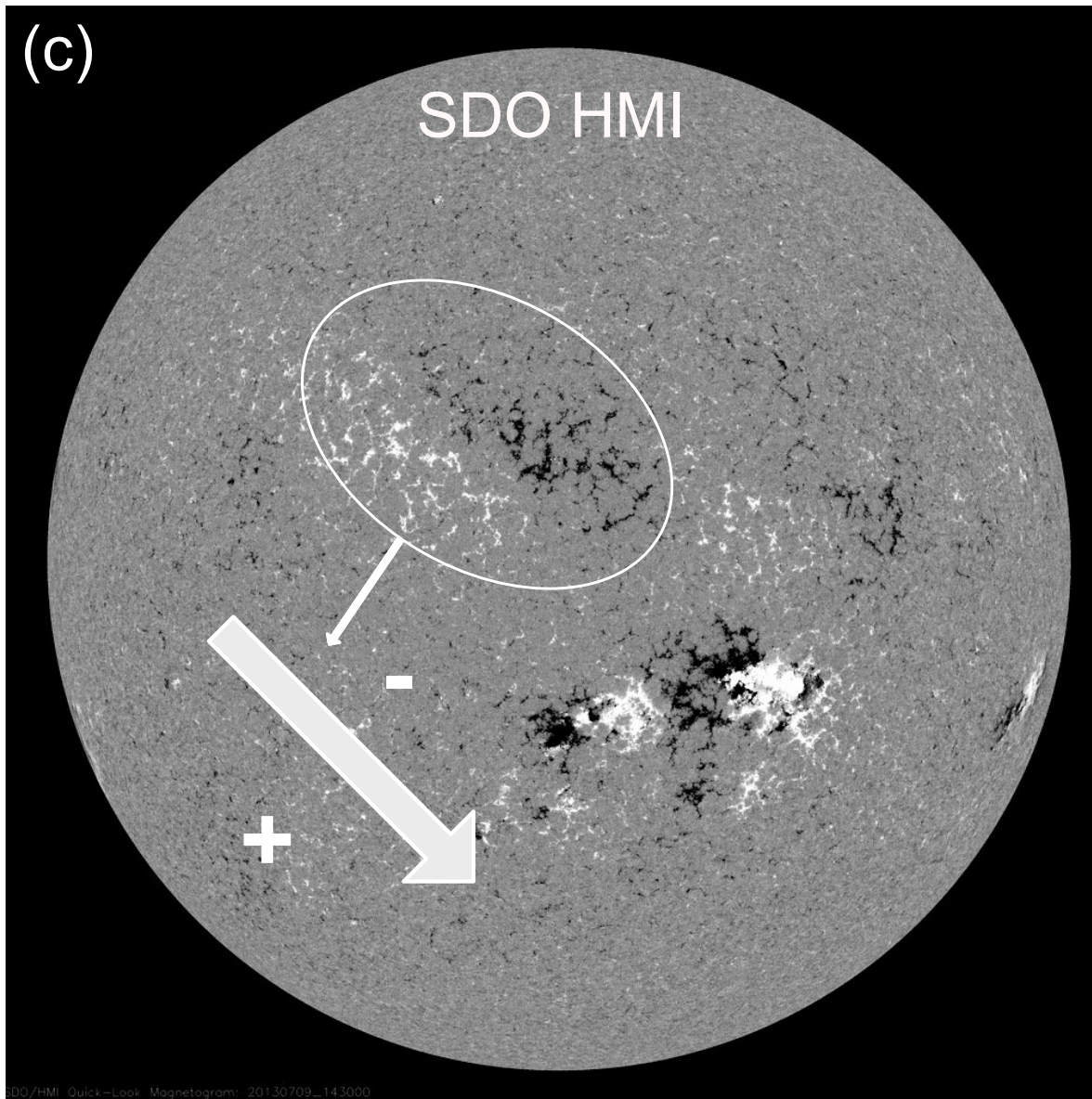
Not included:

- Flattening (next step)
- Flux rope erosion
- CME-CME interaction, CME-CIR interaction
- Manchester, Kilpua, Liu, Lugaz, Riley, Török, Vrsnak 2017 Space Sci. Rev.



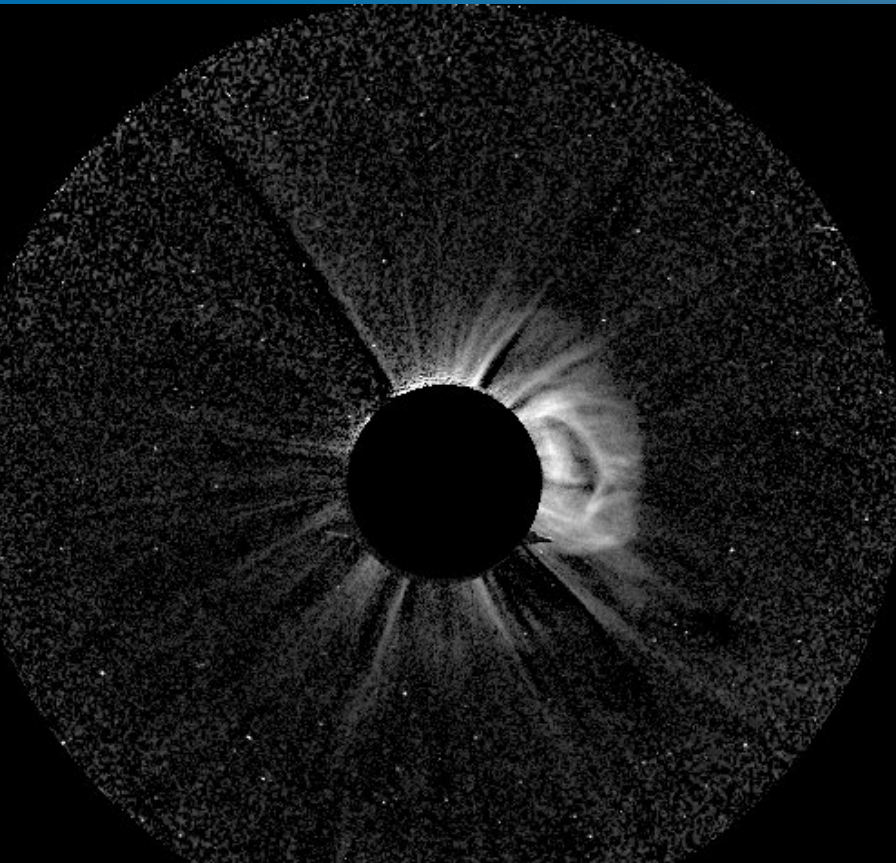
Solar Source

- CME event July 9–13 2013
- Möstl et al. 2018, Space Weather, open access/arXiv

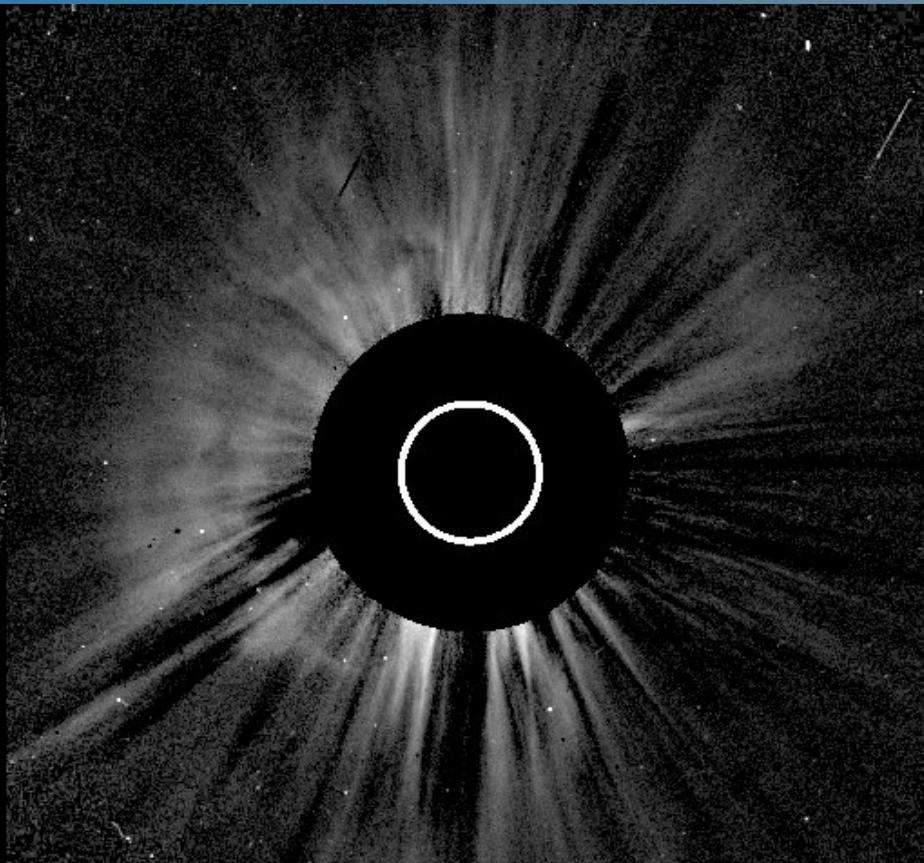


- Palmerio et al. 2018 Space Weather/arXiv
- magnetic structure, left handed, axial field to southwest (follows northern hemisphere rule)

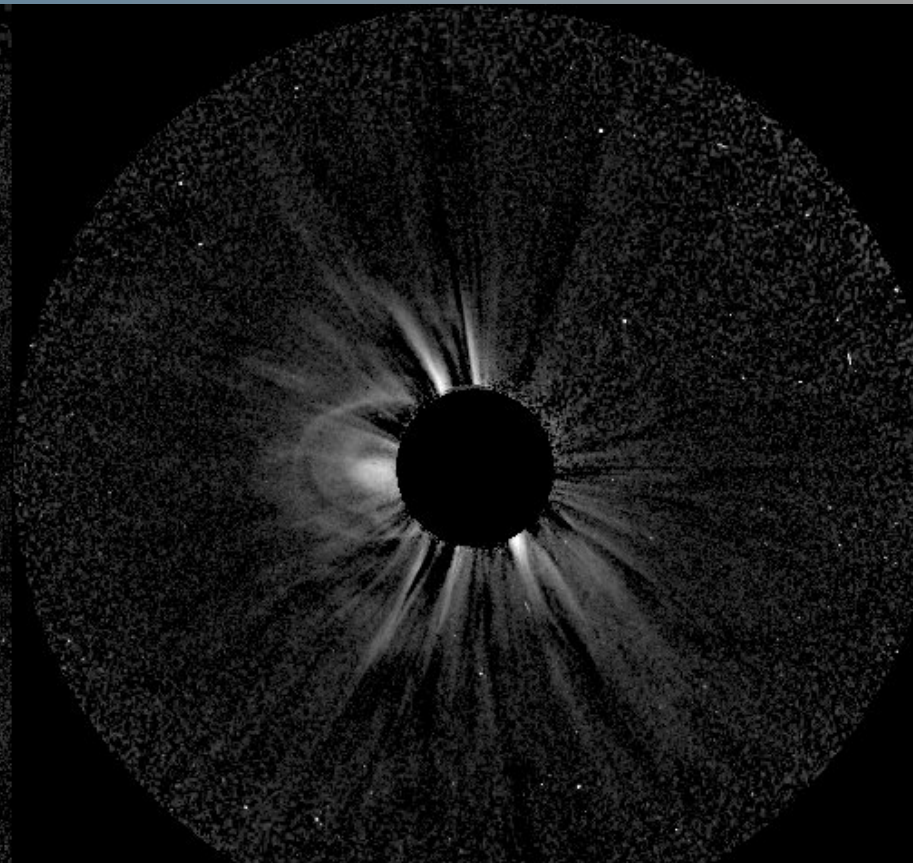
GCS



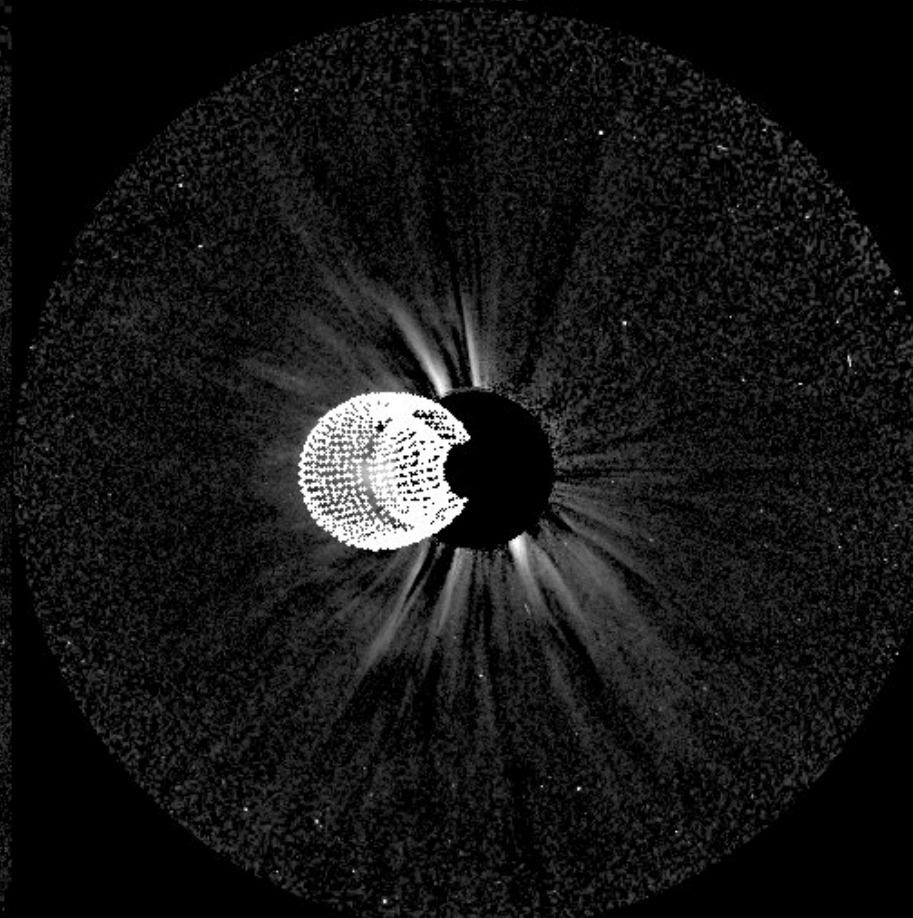
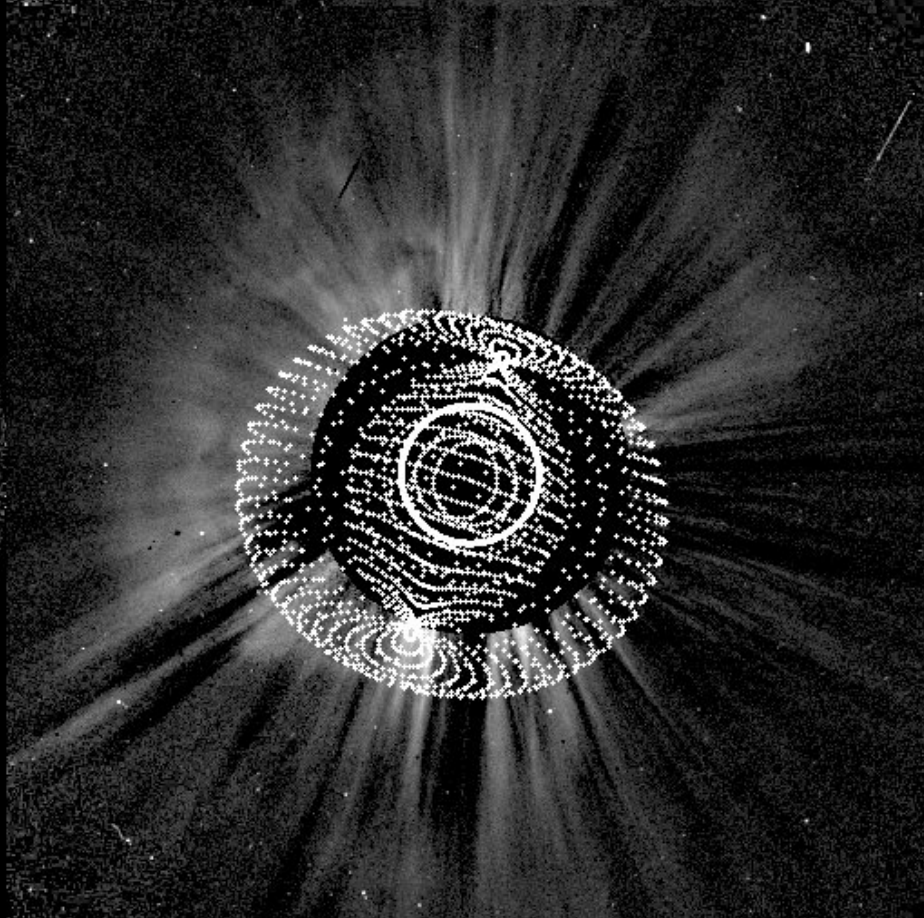
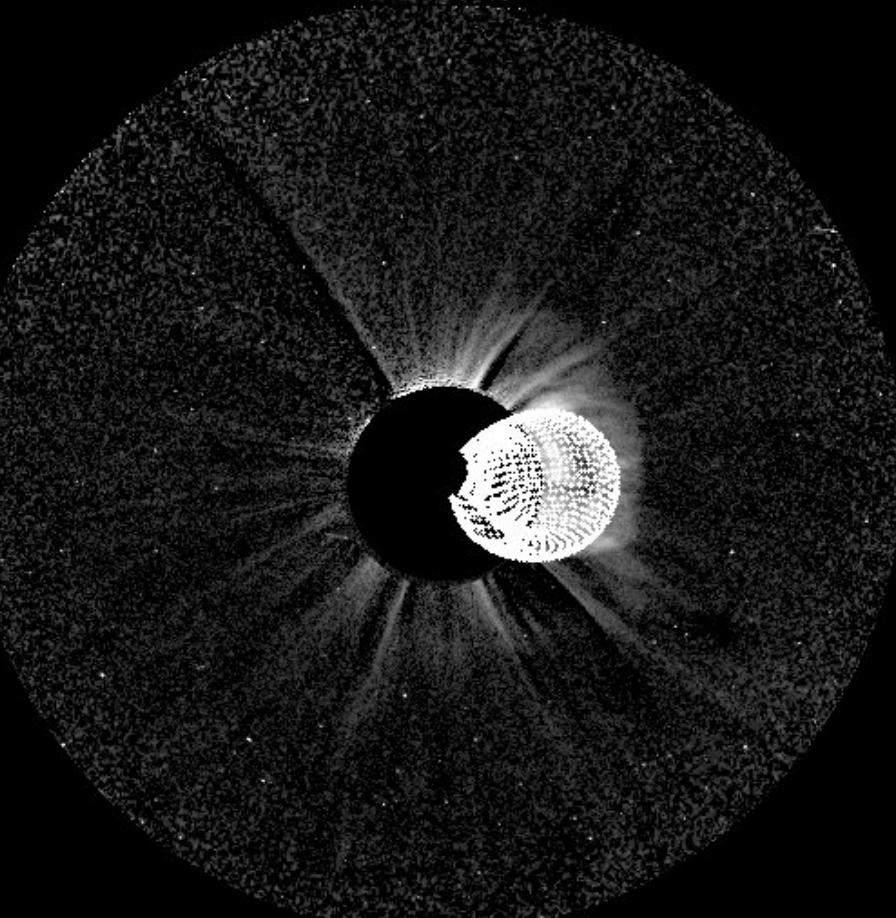
20130709_165400_14c2B.fts



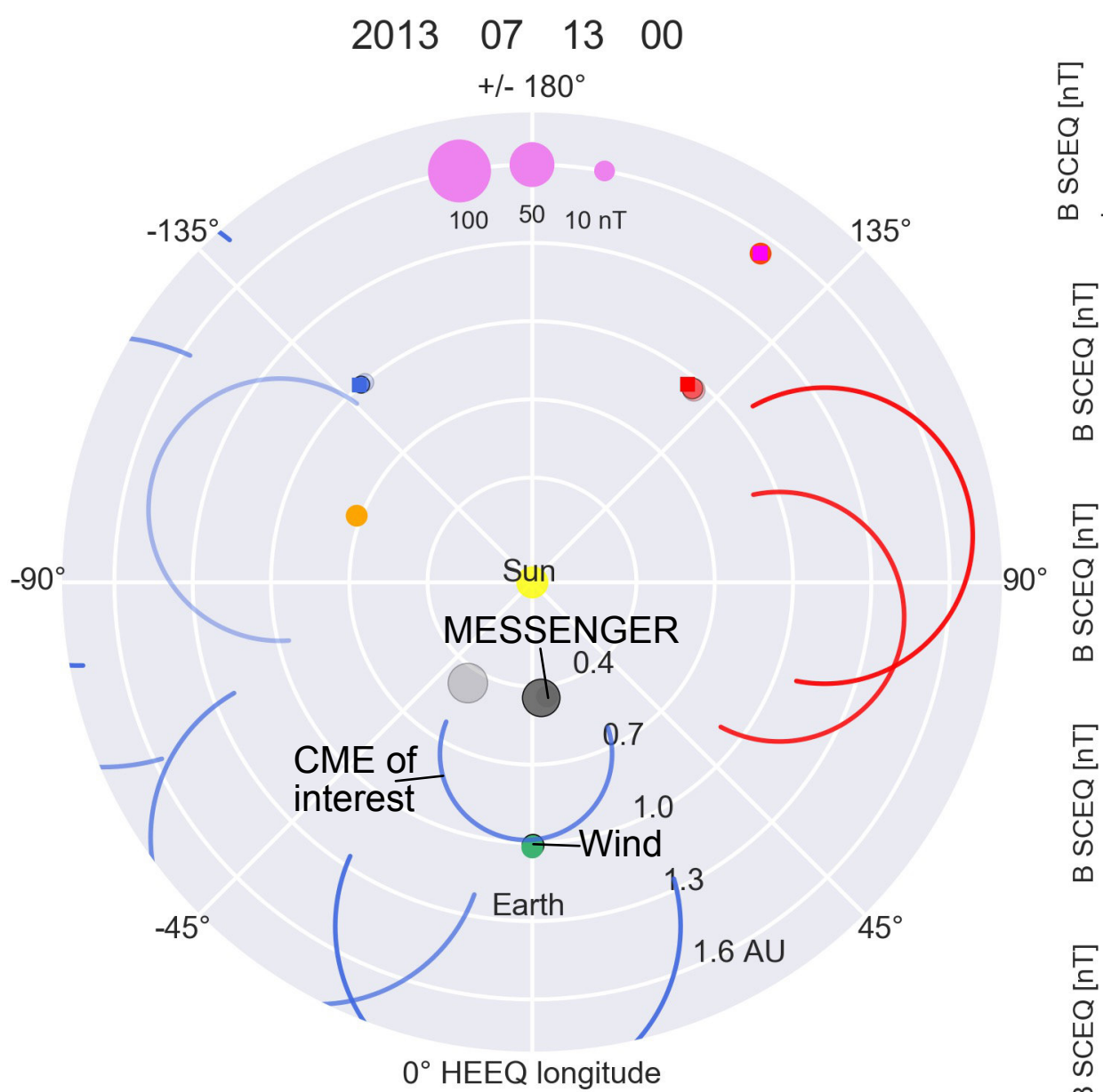
LASCO C2: 2013/07/09 16:58:41



20130709_165400_14c2A.fts



Heliosphere

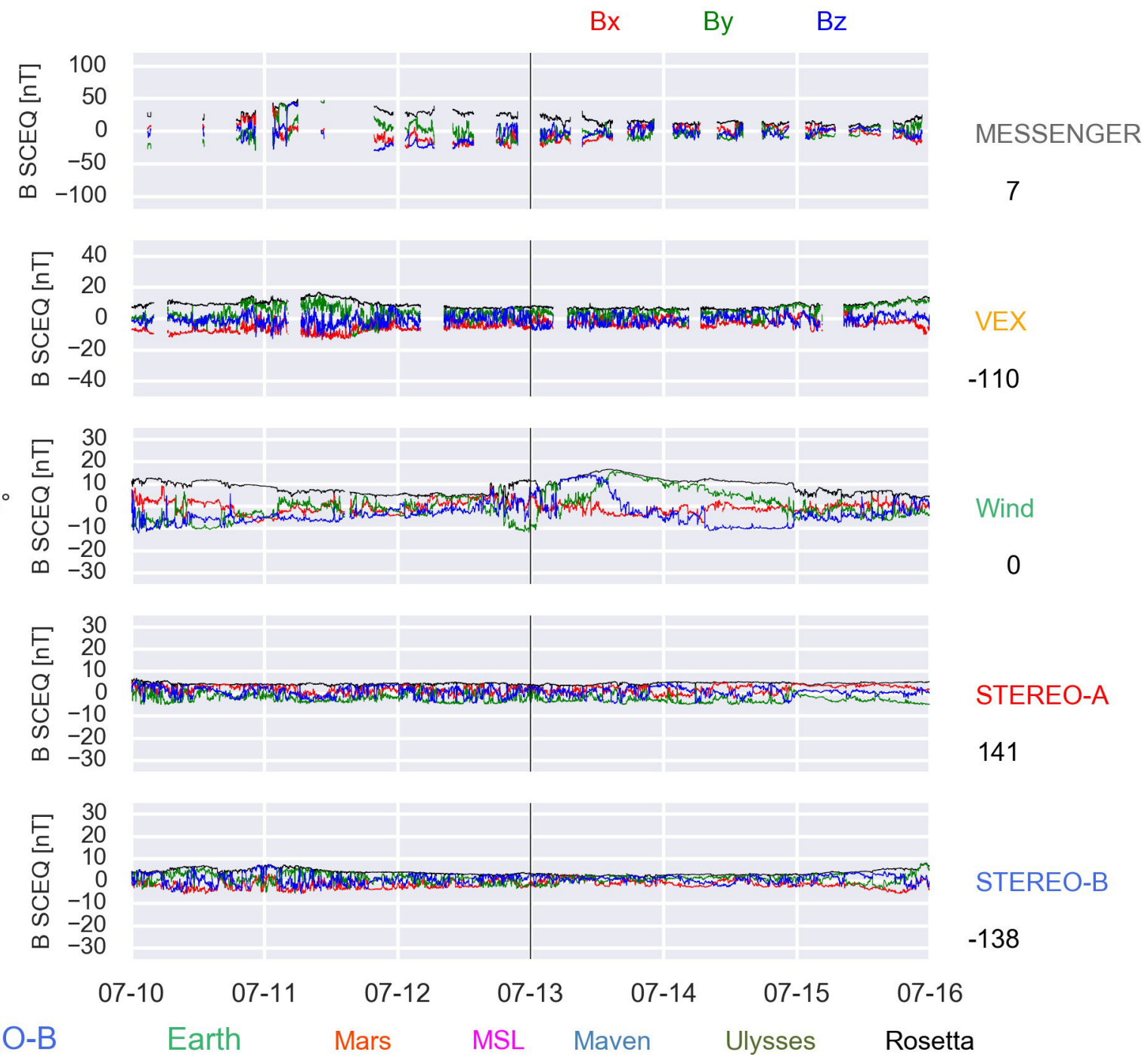


MESSENGER

Venus

STEREO-A

STEREO-B



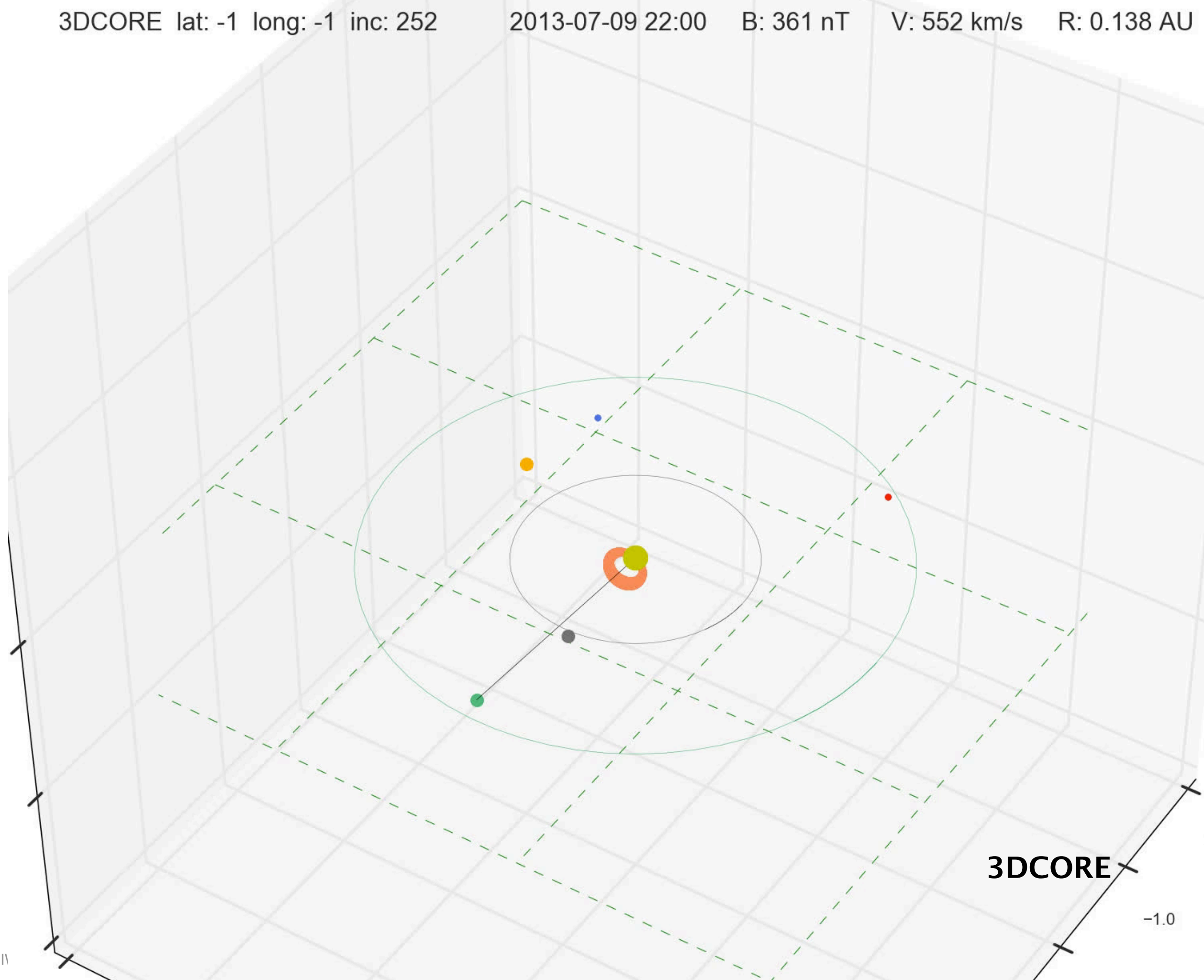
3DCORE lat: -1 long: -1 inc: 252

2013-07-09 22:00

B: 361 nT

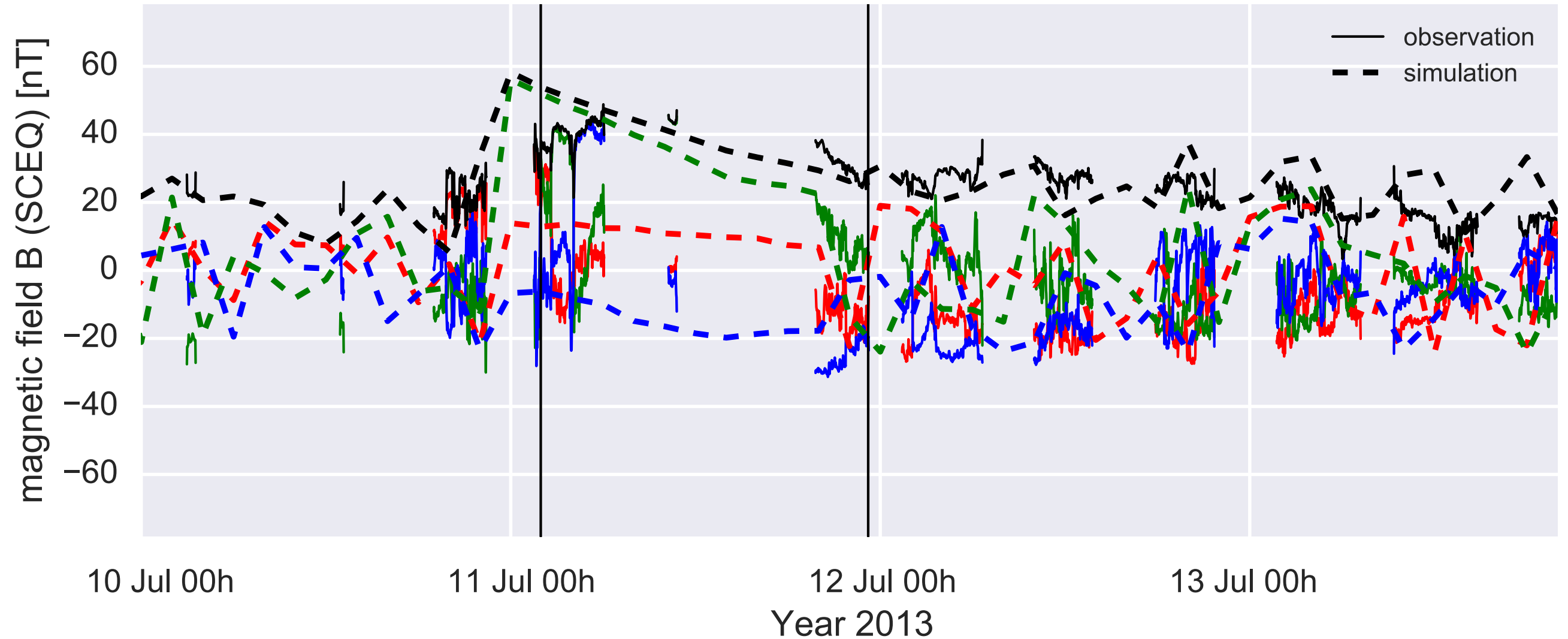
V: 552 km/s

R: 0.138 AU



MESSENGER

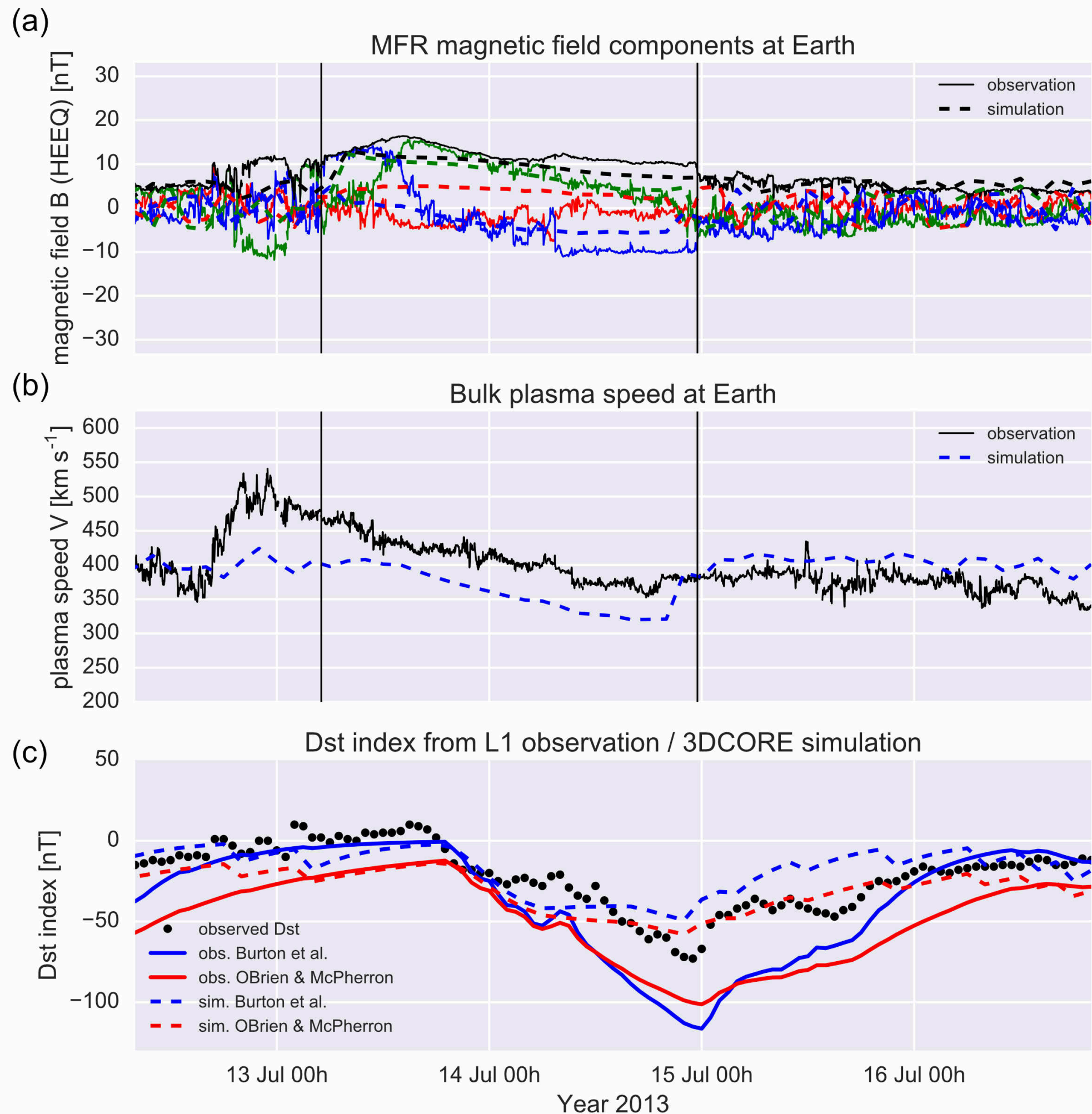
MFR magnetic field components at MESSENGER



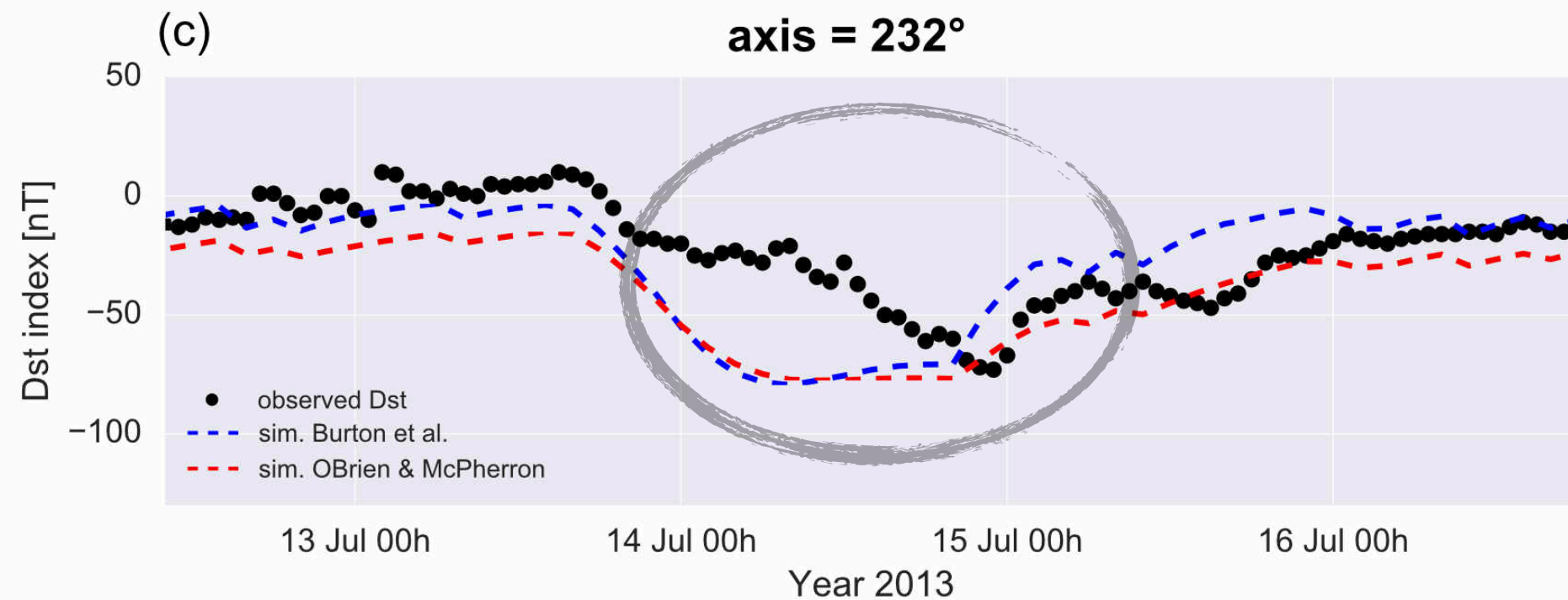
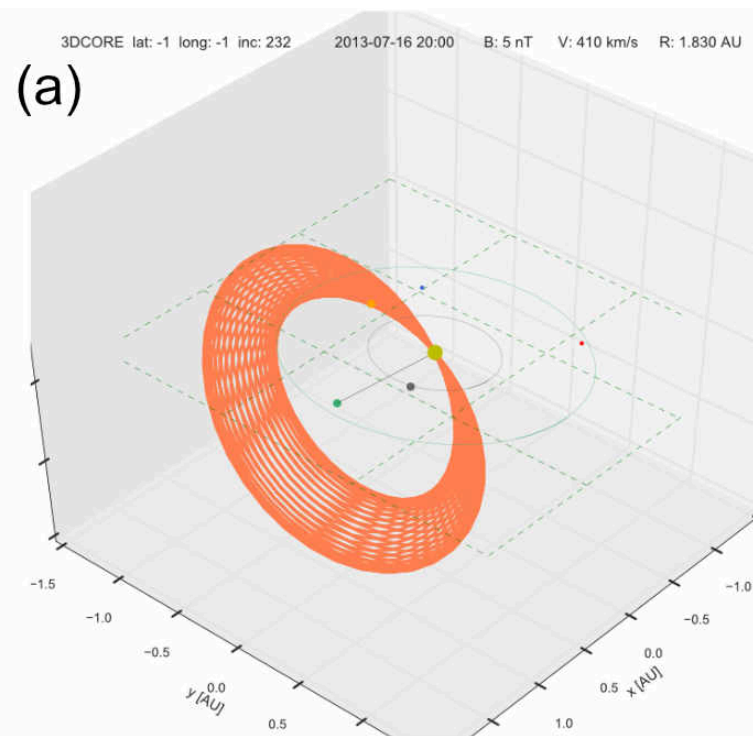
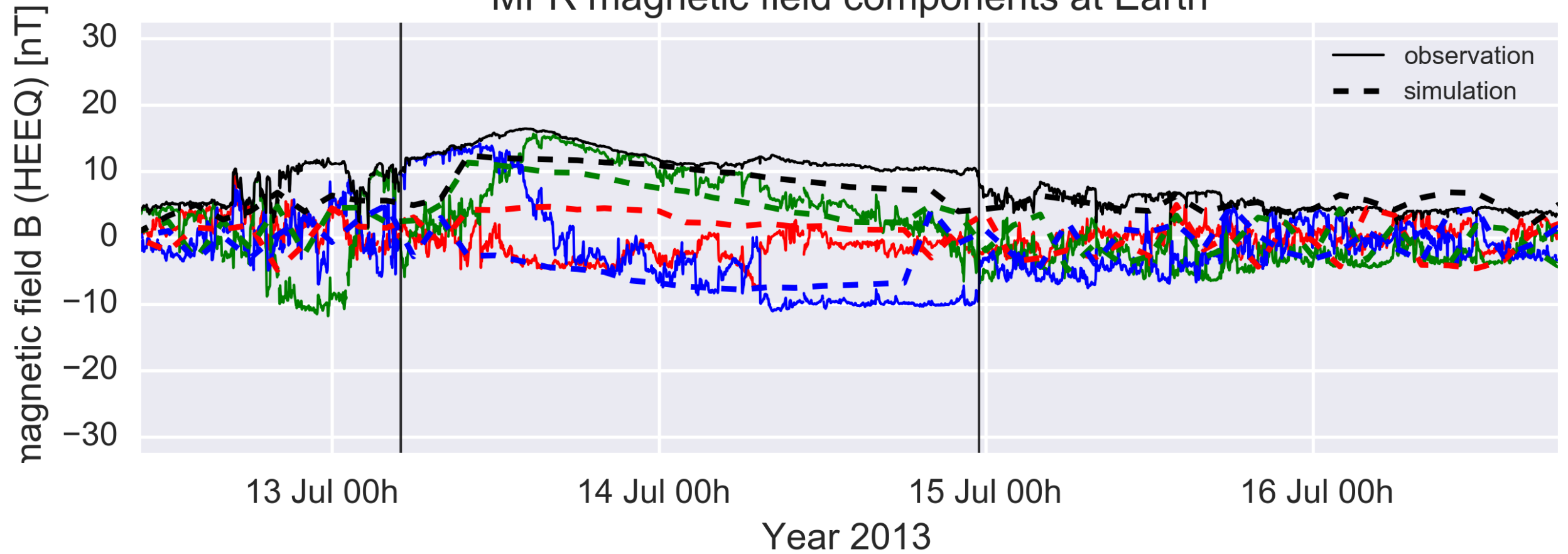
?? drag parameter, initial values: for axial field B_0 , diameter

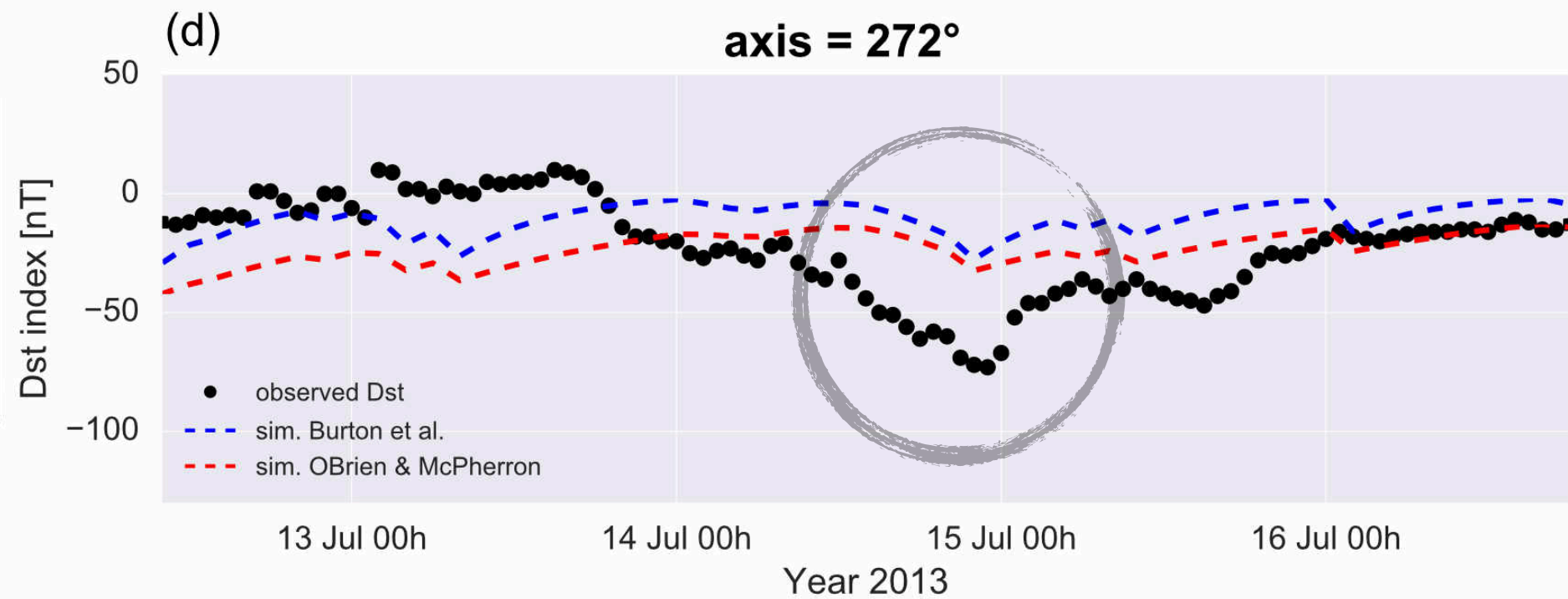
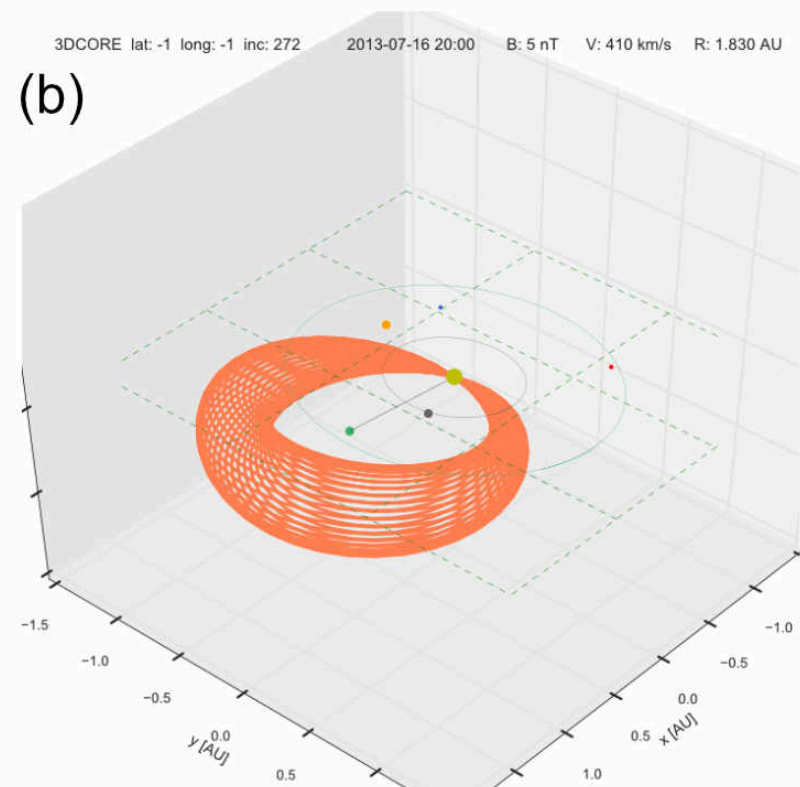
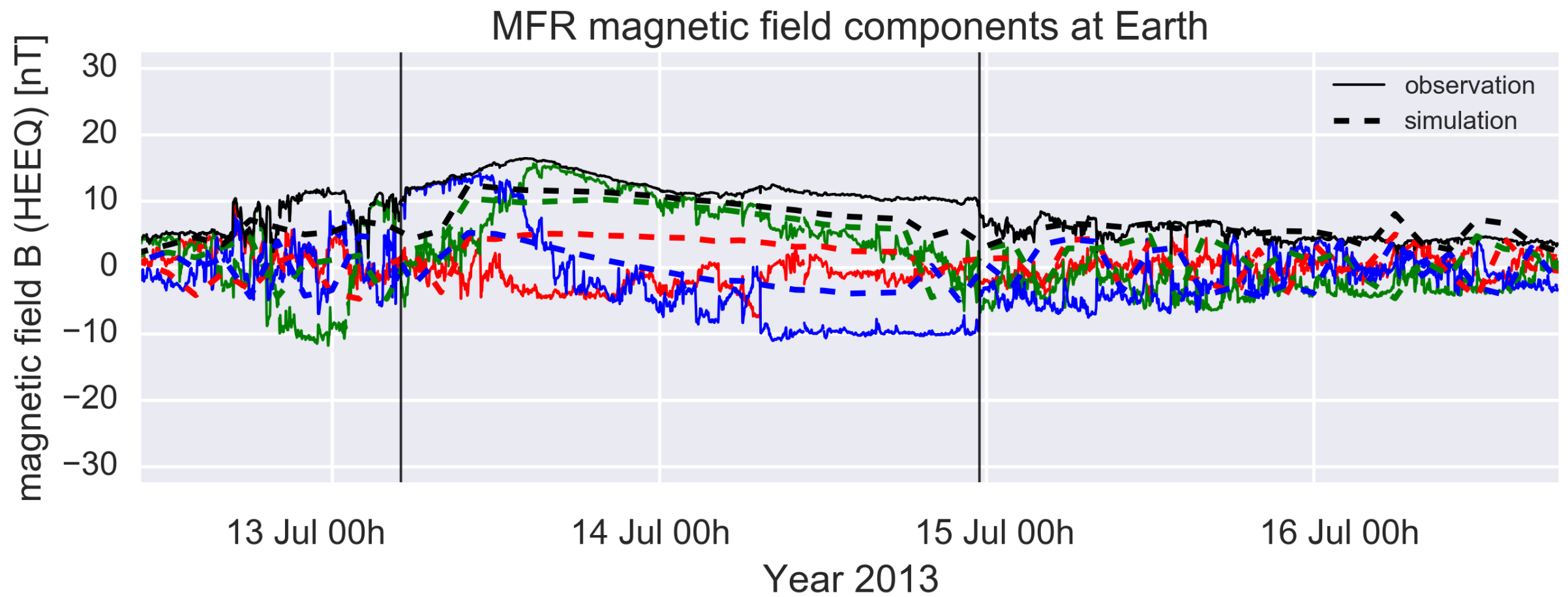
Wind at L1

input:
solar EUV,
magnetogram,
coronagraph GCS
modeling
MESSENGER rope
arrival time,
duration
and maximum
field



MFR magnetic field components at Earth



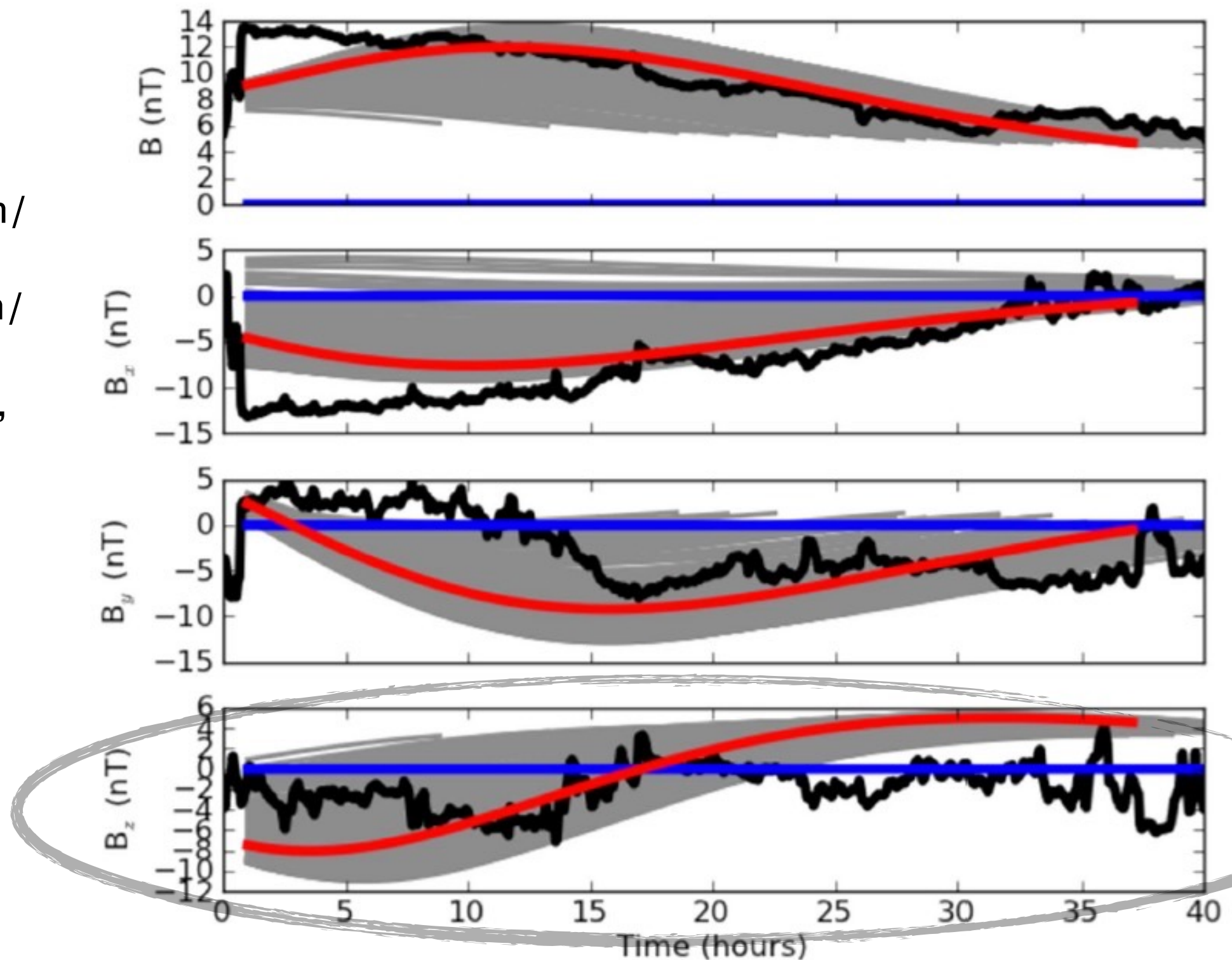


Möstl et al. 2018, Space Weather

FIDO

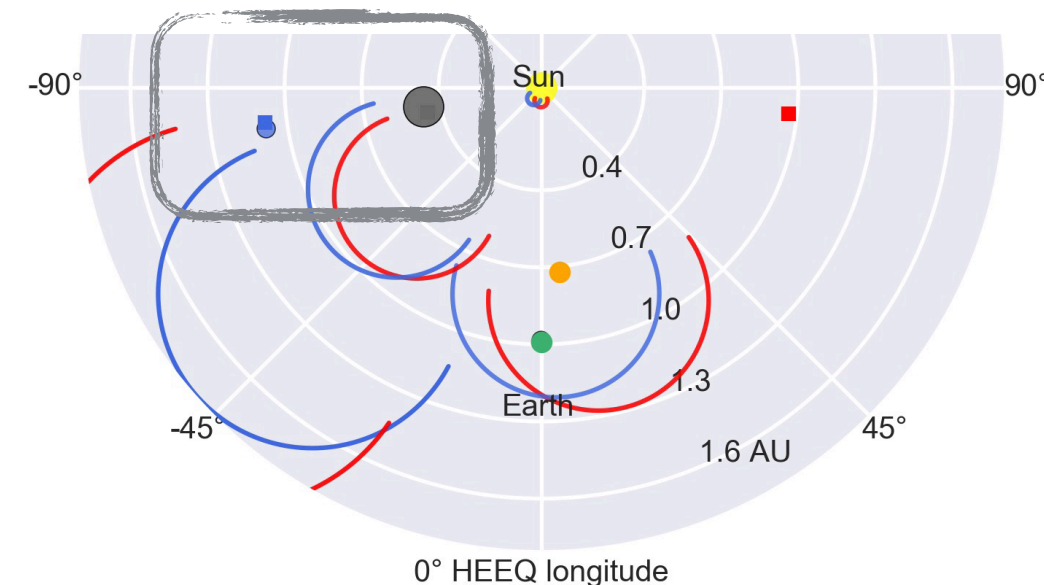
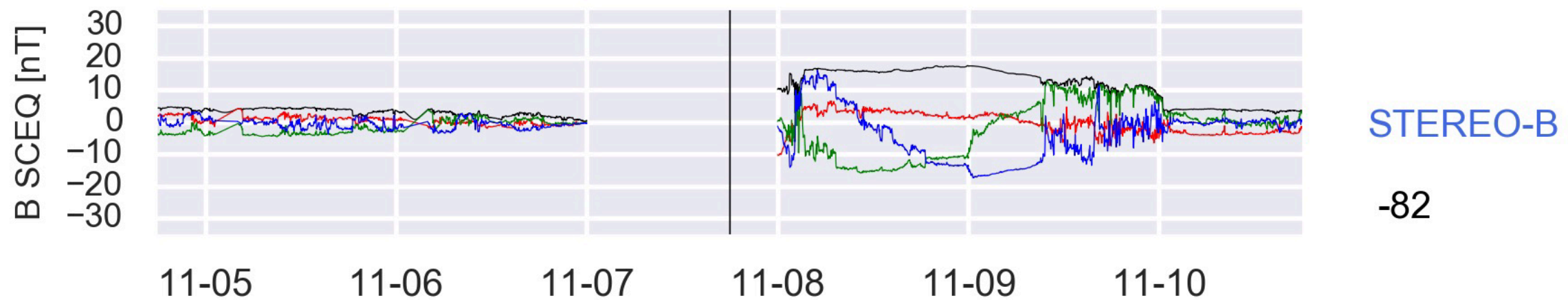
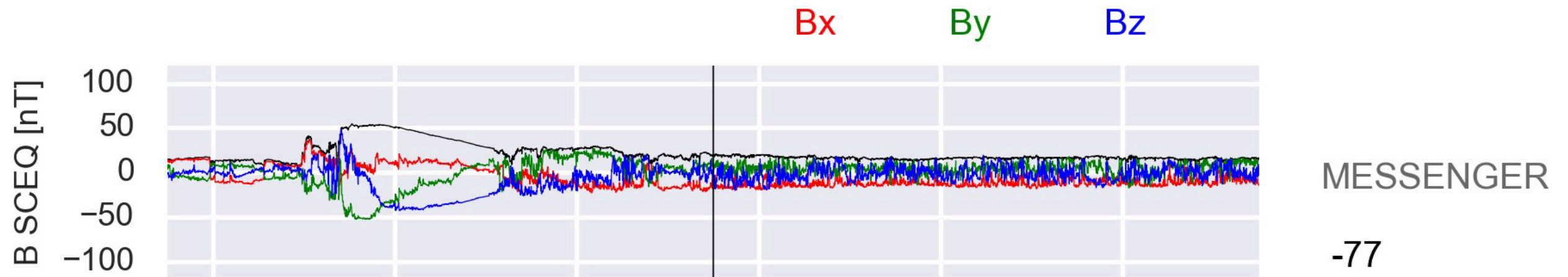
Kay et al. 2017 ApJ

- **black**: ACE
- **red**: FIDO and rotation/deflection
- **blue**: FIDO no rotation/deflection
- **gray**: 150 simulations, variation:
 - latitude $\pm 5^\circ$
 - longitude $\pm 10^\circ$
 - tilt $\pm 10^\circ$
- CME flux rope field profile at L1 is highly sensitive to the CME initial parameters



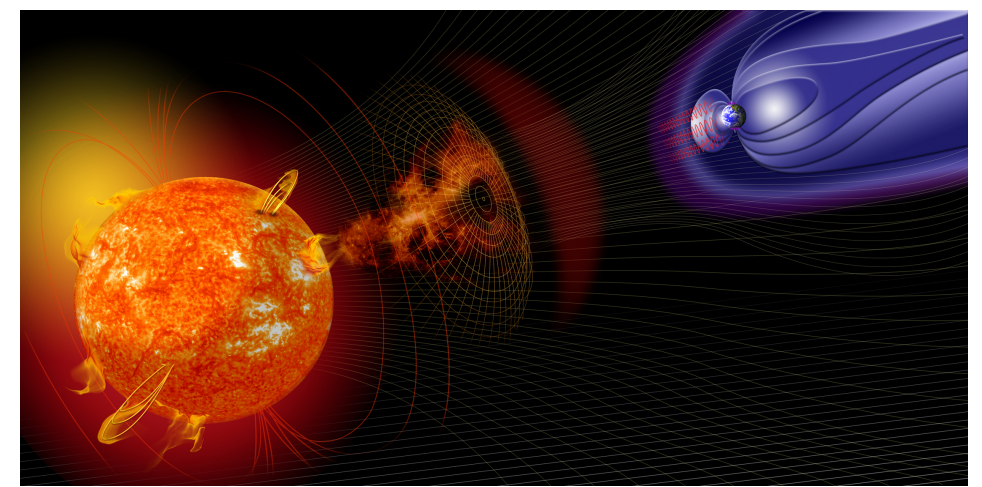
2011 February 15 CME

Future work: more events



- 5 November 2010
Good et al. 2018 Sol. Phys.,
Amerstorfer et al. 2018 Space Weather (revised)
- Catalogs, movies all available on my **figshare** site and **helcats-fp7.eu**
- movies on youtube

3DCORE take-away



- **3DCORE is open source in python** figshare/github ([Möstl et al. 2018](#))
- More events and sensitivity analyses needed ([Kay et al. 2017](#))
- Find constraints on free parameters by combination with solar modeling, coronagraph/heliospheric imaging, ground based observations, in situ observations < 1 AU, ...
E.g. an **arrival time** e.g. at Mercury ([Orsini et al. 2018](#)) near Sun–Earth line is already a good **constraint on the kinematics** ([Kubicka et al. 2016](#))
- Data from **Parker Solar Probe** close to the Sun may lead to breakthroughs on the 3D CME structure with the help of tools like 3DCORE